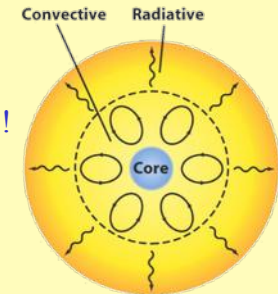


Stellar Evolution: High Mass Stars



High Mass Stars

- **high mass stars**: $> 8M_{\odot}$
- **theoretical upper limit** of $\sim 200M_{\odot}$
- **fusion** $\sim$ via **CNO cycle**
- **very short main sequence** lifetime
(eg) $10M_{\odot}$: ~ 30 million years!
- **evolution** proceeds as for **intermediate mass stars** until the **end of He fusion**



CLICKER: Why do we think stars cannot form** with masses greater than $\sim 200 M_{\odot}$?

- (a) fusion pressure exceeds gravity in outer layers
- (b) rapid rotation at that mass rips the star apart
- (c) $> 200 M_{\odot}$ stars immediately become black holes
- (d) molecular clouds never contain that much mass

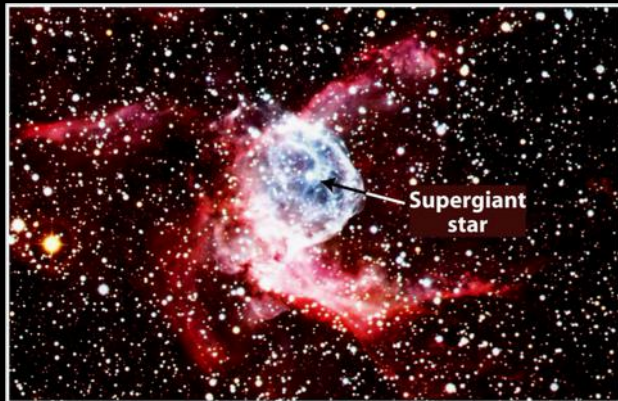
- as core **contracts & heats up** > 100 million K, **star** moves beyond **He burning**
- **envelope greatly** expands: **Supergiant**

(eg) **Betelgeuse in Orion, Deneb in Cygnus**



- undergo **tremendous mass loss** from **envelope**
- **transient** periods of **pulsation** (**Cepheids**)

Mass Loss ($40 M_{\odot}$, *Canis Major*)



- **core** sequentially fuses *more massive nuclei*:



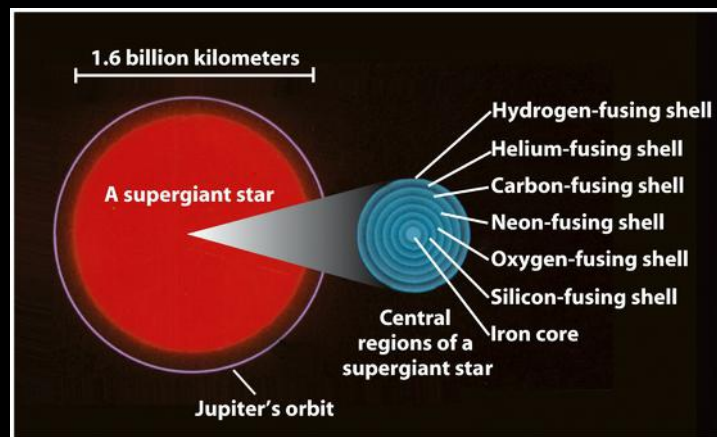
- **fusion reactions** can proceed *in multiple ways*
(eg) $^{16}O + ^4He \Rightarrow ^{20}Ne$ & $^{12}C + ^{12}C \Rightarrow ^{20}Ne + ^4He$

- other **fusion** "byproducts" include *Na, P, Ni*

- temperatures: *600 million K (C)*, *3 billion K (Si)*

- **fusion** of *Si* creates iron (*Fe*) but *Fe cannot* fuse

Supergiant Core

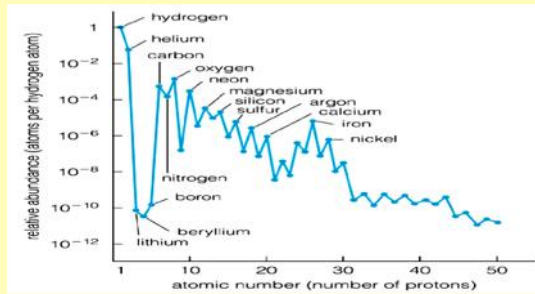


$25M_{\odot}$ Fusion Lifetimes

table 22-1 Evolutionary Stages of a $25M_{\odot}$ Star

Stage	Core temperature (K)	Core density (kg/m ³)	Duration of stage
Hydrogen fusion	4×10^7	5×10^3	7×10^6 years
Helium fusion	2×10^8	7×10^5	7×10^5 years
Carbon fusion	6×10^8	2×10^8	600 years
Neon fusion	1.2×10^9	4×10^9	1 year
Oxygen fusion	1.5×10^9	10^{10}	6 months
Silicon fusion	2.7×10^9	3×10^{10}	1 day

- **Fe** is “*endpoint*” for both *fusion* & *fission* since it has *lowest mass-per-nuclear particle* (*eg*) cannot make Fe nuclei less “massive”



(*eg*) energy released in **H** fusion since $m_{\text{He}} < 4m_{\text{H}}$

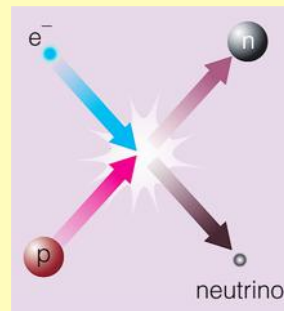
- *peak* in abundance at **Fe** in the *periodic table*

- *core* is $> 1.4 M_{\odot}$ and $\sim$ size of Earth

Q: What happens when core is entirely made of Fe?

- *core collapses* in a *fraction of a second*
- electrons *already* crushed together
- *electron degeneracy ineffective* (**Chandrasekhar**)
- *no further fusion* possible
- *temperatures* rise to over 5 billion K

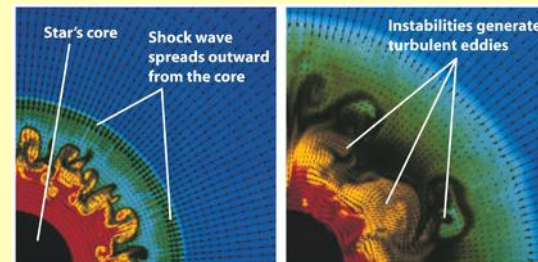
- *protons & electrons* fuse into *neutrons*
- *neutron degeneracy* opposes collapse
- *core* is size of *city* (~ 20 km)
- 1 *tsp* $\sim$ mass of a *mountain*



(*eg*) core density $\sim 10^{17} \text{ kg/m}^3$ (like *atomic nuclei*)

- *extremely rigid* $\Rightarrow$ *in-falling* material “*bounces*”
- forms an outward moving *shockwave* (10^4 km/s!)

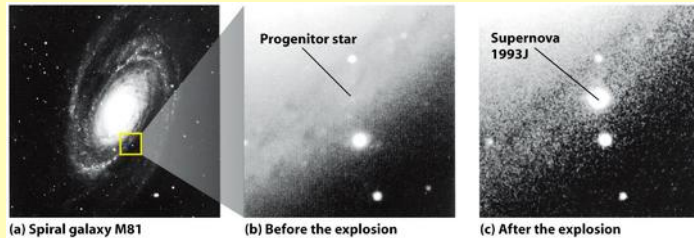
- rest of *envelope* rapidly falls inward ($\sim 0.1c$)
- struck by outgoing *shockwave/neutrinos*
- *envelope* is torn apart, *compressing/heating* it & *creating some heavy elements* on *periodic table*



“Not only are we a part of the Universe... but the universe is a part of us.” Neil deGrasse Tyson

(Type II) Supernovae

- **luminosity** of explosion can reach $10^9 L_{\odot}$
- **energy released in a few seconds** is **100x that produced by the Sun during its entire lifetime**



- for weeks, **supernovae** shine like *a small galaxy*

Historical Supernovae

- only 5 **naked eye supernovae** in last 1000 years:
1006, 1054, 1181, 1572, 1604
- **all of these took place within our galaxy**
- **1006**: bright enough to cast shadows at night
- **1054**: bright enough to **read by** at night
- **1572**: Tycho Brahe's supernova
- 1987 detection of energetic **neutrinos**...

Supernova 1987a (LMC, ~200,000 ly)



Supernova Remnants (LMC)

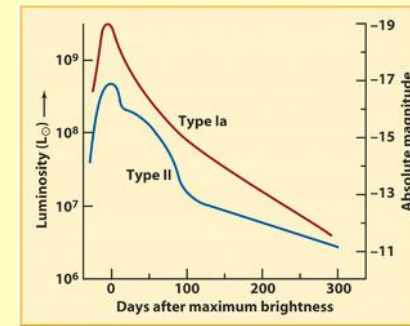


- filtered view with stars “subtracted”
- H (red), O (green), S (yellow)

Supernova Remnant c. 1054 (Crab Nebula - Taurus)



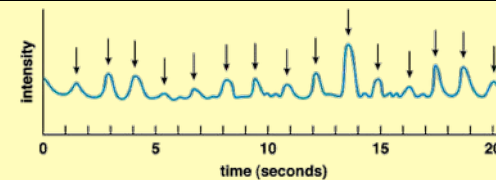
Light Curves



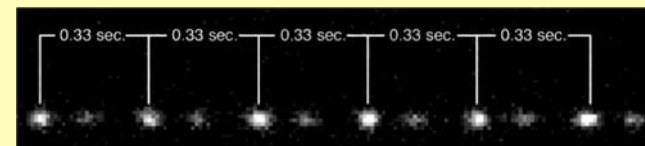
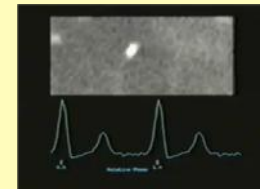
- **supernovae** yield *distances*, since *peak luminosity* (how *bright* they *should* be) is *known* (variations?)

Neutron Stars

- **neutron** discovered in 1932
- **Fritz Zwicky** at *CalTech* proposed **Neutron Star**
- **neutron star**: remnant **core** of star *after supernova*
- **electrons & protons** combine to form **neutrons**
- **core** supported by **neutron degeneracy**
- $\sim 3 M_{\odot}$ **neutron star** has **escape velocity** of $0.5 c$
- seemed *too outrageous* to possibly be *true*



- 1967: **pulsating radio signals** detected in **Cygnus** by grad student Jocelyn Bell
- **LGM's** ("**Little Green Men**")
- too *common & regular*: renamed **pulsars**



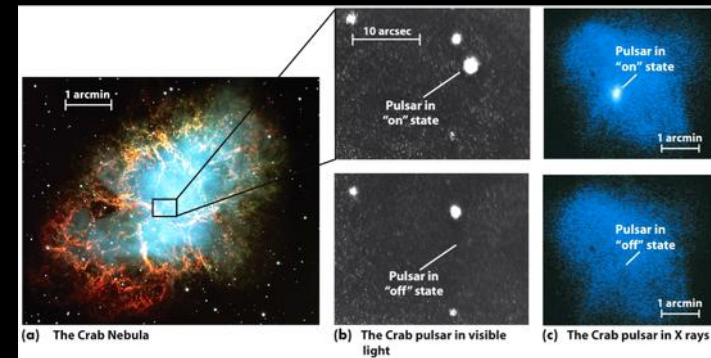
- **pulsars** found at centers of *Crab, Vela supernovae*

Q: What could cause the pulsing signals?

- **pulsars spin rapidly**
- up to 1000x per sec!
- must be *small, dense*
- **intense magnetic field** (10^{12} G) expels radiation in "**beams**" - like *lighthouse*

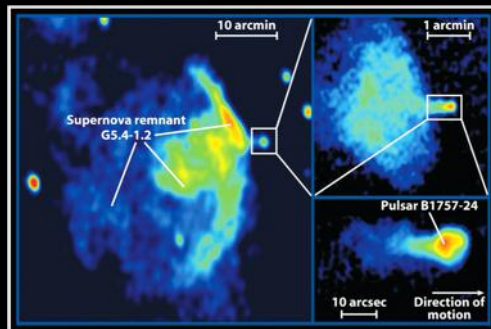


Crab Nebula Pulsar



Runaway Neutron Stars?

- *supernova* may explode *asymmetrically*, *ejecting* a *neutron star* at *high speeds* (>1000 km/s ?)



Neutron Star Limit

- **gravity** balanced by **neutron degeneracy** & **strong nuclear force**

(eg) **strong force** binds **neutrons** and **protons**; **strongly resists compression** - like **billiard balls**

- **Neutron Star Limit** not as well known as **Chandrasekhar Limit** (**white dwarfs**), since **strong force** is less well understood
- if **core** $>3M_{\odot}$, **no force known can stabilize core**

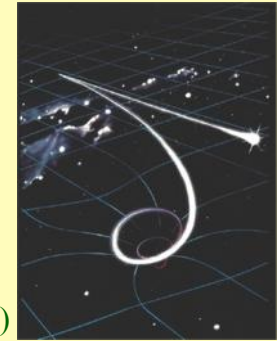
CLICKER: A neutron star is

- (a) left behind after a Type Ia supernovae explosion
 (b) created immediately after fusion of H ceases
 (c) one possible remnant of Type II supernovae
 (d) at the center of a planetary nebulae

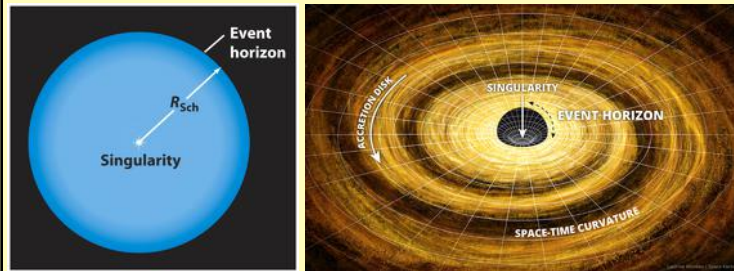
- **core** collapses to a **point** (“singularity”)
- “**hole**” in the fabric of “**spacetime**”
- **gravity** so powerful that **escape speed exceeds c**
- even **light cannot** escape!

“Black Holes”

- proposed in 1700s by **Laplace** using **Newton's escape speed**
- coined by **John Wheeler** (1967)



Features of Black Holes



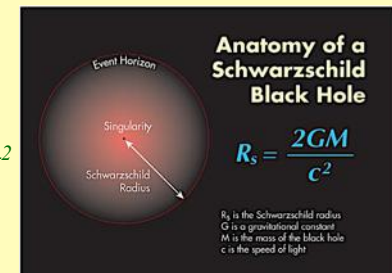
- **event horizon**: visible **boundary** of a **black hole**
- at **event horizon**, the **escape velocity** = **c**
- an **accretion disk** of **hot, rotating gas** may exist

Schwarzschild Radius

- **R_{sch}** is distance from **singularity** to **event horizon**

$$R_{sch} = 2GM/c^2$$

- **$G = 6.67 \times 10^{-11} \text{ Nm}^2/\text{kg}^2$**
- **$c = 3.0 \times 10^8 \text{ m/s}$**
- **M in kg, R in m**



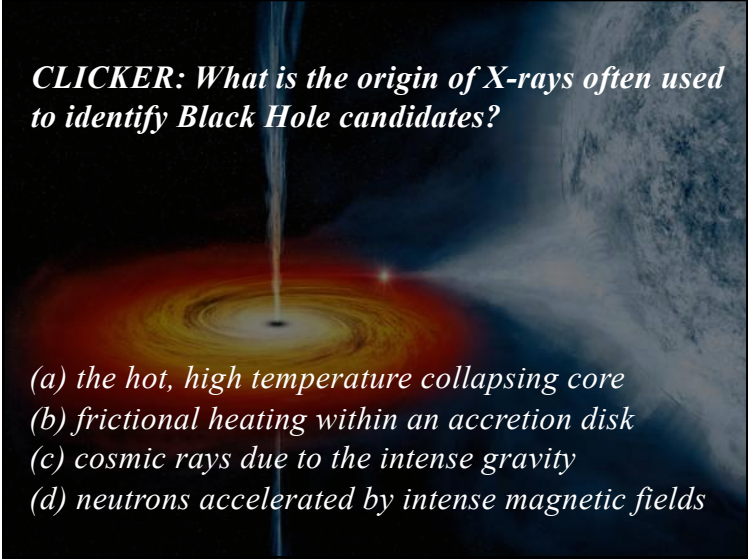
(eg) A $5 M_{\odot}$ star forms a black hole of $R_{sch} = 15 \text{ km}$

Finding Black Holes

Q: How do you look for “nothing”?

- **accretion disk** emissions
- **gravitational interactions** with other objects
- **gravitational lensing** by the black hole
- **observational uncertainties** can make it difficult to differentiate **black holes** from **neutrons stars**

CLICKER: What is the origin of X-rays often used to identify Black Hole candidates?

- 
- (a) the hot, high temperature collapsing core
 - (b) frictional heating within an accretion disk
 - (c) cosmic rays due to the intense gravity
 - (d) neutrons accelerated by intense magnetic fields

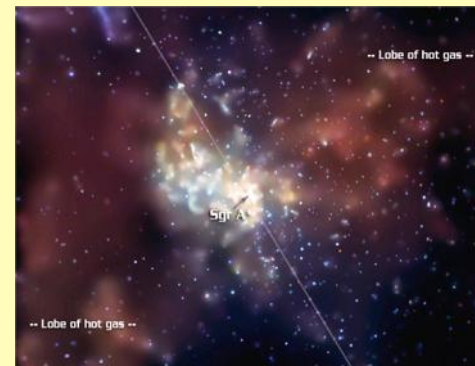
Black Hole candidate Cygnus X-1

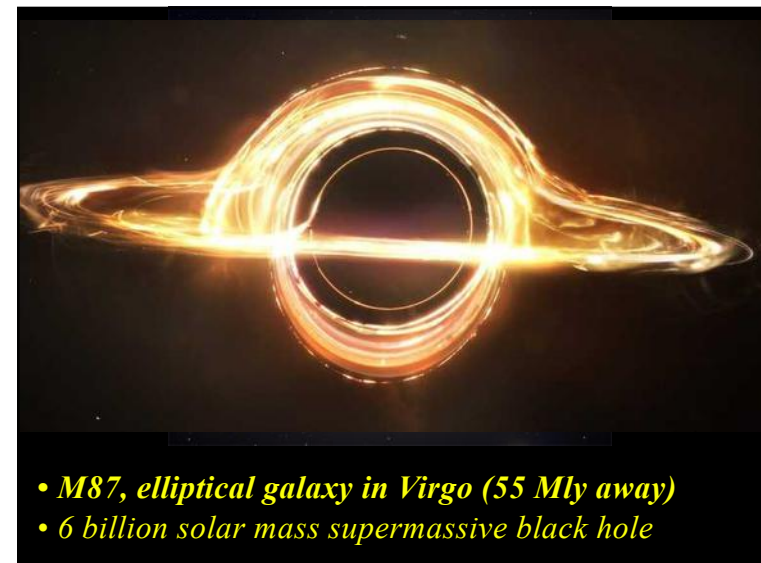
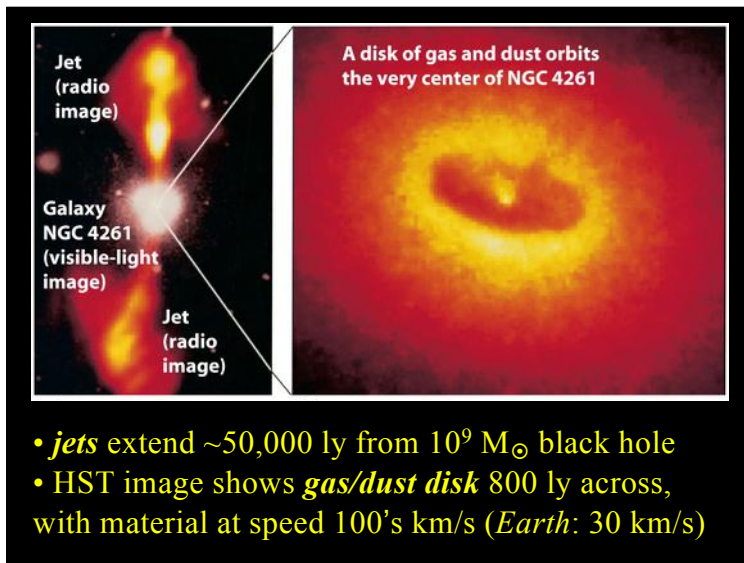


- discovered in 1970 by *Uhuru X-Ray satellite*
- $\sim 30 M_{\odot}$ B01 star & $\sim 15 M_{\odot}$ object co-orbiting

Supermassive Black Holes

- **radio & X-ray measurements** of the centers of galaxies show many have **strong** emission sources





Black Hole FAQ



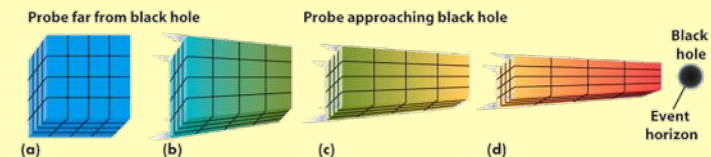
*Q: What would happen if **Sun** became a BH?*

Q: Are BH's giant interstellar "vacuums"?

Q: Do BH's last forever?

Falling into a stellar BH

- *time & space distort close to the **event horizon***



- wavelengths *lengthen* (*gravitational redshift*)
- *cannot* (?) survive journey to the *singularity*

